

# Discovery of a supercluster in the Zone of Avoidance in Vela

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## ABSTRACT

We report the discovery of a potentially major supercluster that extends across the Galactic plane in the constellation of Vela, at a mean recessional velocity of  $\sim 18\,000\text{ km s}^{-1}$ . Recent multiobject spectroscopic observations of this Vela supercluster (VSCL), using AAOmega+2dF and the Southern African Large Telescope, confirm an extended galaxy overdensity in the Zone of Avoidance (ZOA) located where residual bulk flows predict a considerable mass excess. We present a preliminary analysis of  $\sim 4500$  new spectroscopic galaxy redshifts obtained in the ZOA centred on the Vela region ( $l = 272.5 \pm 20^\circ$ ,  $b = 0^\circ \pm 10^\circ$ ). The presently sparsely sampled data set traces an overdensity that covers  $25^\circ$  in Galactic longitude on either side of the Galactic plane, suggesting an extent of  $25 \times 20\text{ deg}^2$ , corresponding to  $\sim 115 \times 90\ h_{70}\text{ Mpc}$  at the supercluster redshift. In redshift space, the overdensity appears to consist of two merging wall-like structures, interspersed with clusters and groups. Both the velocity histogram and the morphology of the multibranching wall structure are consistent with a supercluster classification.  $K_s^0$  galaxy counts show an enhancement of  $\sim 1.2$  over the survey area for galaxies brighter than  $M_K^*$  at the VSCL distance, and a galaxy overdensity of  $\delta = 0.50\text{--}0.77$  within a photometric redshift shell around the VSCL, when compared with various Two Micron All-Sky Survey samples. Taking account of selection effects, the VSCL is estimated to contribute  $v_{\text{LG}} \gtrsim 50\text{ km s}^{-1}$  to the motion of the Local Group.

**Key words:** techniques: spectroscopic – surveys – galaxies: clusters: general – galaxies: distances and redshifts – cosmology: observations – large-scale structure of Universe.

## 1 INTRODUCTION

Galaxies are distributed in a cosmic web consisting of high-density clusters and galaxy groups embedded in walls and filaments surrounding large low-density regions (e.g. Jones et al. 2009; Tempel et al. 2014). The largest agglomerations, superclusters, can contain dozens of massive galaxy clusters (Proust et al. 2006; Chon, Böhringer & Nowak 2013; Einasto et al. 2014). Such galaxy overdensities exert gravitational perturbations on the smoothly expanding Universe and induce coherent flows over large volumes, as, for instance, the Local Group (LG) motion imprinted as a dipole in the cosmic microwave background ( $\sim 622\text{ km s}^{-1}$ ; Fixsen et al. 1996).

Despite numerous studies (e.g. Hudson et al. 2004; Erdođdu et al. 2006; Kocevski & Ebeling 2006; Watkins, Feldman & Hudson 2009; Lavaux et al. 2010; Bilicki et al. 2011; Branchini, Davis & Nusser 2012; Hoffman, Courtois & Tully 2015; Springob et al. 2016), the resulting direction and amplitude of the LG peculiar velocity and of local bulk flows remain controversial. The persistent discrepancies may originate, in large part, from the incomplete mapping of large-scale structures in the Zone of Avoidance (ZOA; Kraan-Korteweg & Lahav 2000; Loeb & Narayan 2008), which are excluded in ‘whole-sky’ galaxy surveys, and observed with lower sensitivity at the relevant higher distance range ( $cz \gtrsim 16\,000\text{ km s}^{-1}$ ) in targeted surveys.

To reconcile the discrepancies, a hidden mass overdensity was postulated behind the southern ZOA (Loeb & Narayan 2008). One possibility, a supercluster at or beyond the Shapley concentration (SSC), with  $cz > 16\,000\text{ km s}^{-1}$ , was considered unlikely even if it resided in the ZOA, as there had been no indication of its existence

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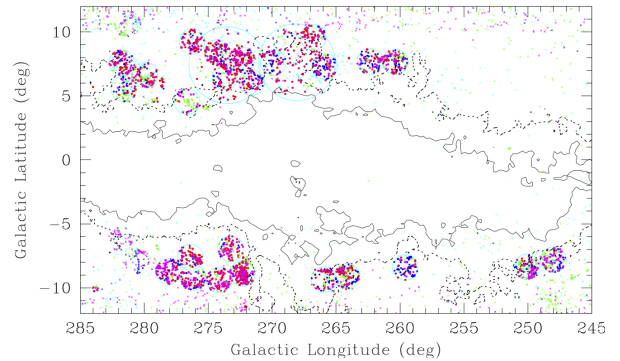
in whole-sky surveys. Since then, however, various new studies implied a considerable mass excess close to the ZOA in the direction of the Vela constellation at higher redshifts (e.g. Hudson et al. 2004; Feldman, Watkins & Hudson 2010; Nusser & Davis 2011; Carrick et al. 2015; Hoffman et al. 2015). And the recent 6dFGS and 2MTF peculiar velocity analyses (Springob et al. 2014; Scrimgeour et al. 2016) require a residual bulk flow ( $273 \text{ km s}^{-1}$ ) arising from that same general direction, which is generated beyond a distance of  $230 h_{70} \text{ Mpc}$ . However, at these depths, the Vela area is not mapped by current spectroscopic surveys, which are too shallow or do not sample  $|b| < 10^\circ$  (e.g. 2MASS Redshift Survey, Huchra et al. 2012, or 2M++, Lavaux & Hudson 2011).

We have embarked on a long-term programme to map the large-scale galaxy distribution behind the Milky Way, giving particular emphasis to the Vela ZOA area ( $l = 272.5 \pm 20^\circ$ ,  $b = 0^\circ \pm 10^\circ$ ). Our earliest ZOA multiobject spectroscopic data close to Vela revealed clear hints of a substantial galaxy overdensity at  $18\,000 \text{ km s}^{-1}$  (see figs 3 and 5 in Kraan-Korteweg et al. 1994) based on observations made with Optopus on the ESO 3.6-m telescope in the Hydra/Antlia ZOA galaxy survey ( $295^\circ \gtrsim l \gtrsim 275^\circ$ ; Kraan-Korteweg 2000). Speculations about a possible connection across the ZOA between the Horologium-Reticulum and Shapley superclusters were alluded to. Subsequent observations of galaxies in the adjacent Vela ZOA region ( $275^\circ \gtrsim l \gtrsim 250^\circ$ ; Kraan-Korteweg & Lahav 2000) with the 6dF multifibre spectrograph, on the UK Schmidt Telescope, of two observed contiguous ZOA fields (350 redshifts) revealed a highly significant peak at the same redshift. The combined data set implies a galaxy overdensity of surprising extent for its mean redshift. An analysis of the extinction-corrected magnitudes of galaxies in the deep optical Vela ZOA catalogue suggested further clustering at this approximate distance range, and also on the other side of the Galactic plane, linking up to two X-ray clusters (CIZA J0812.5–5714 and CIZA J0820.9–5704; Kocevski & Ebeling 2006) at a similar redshift ( $\sim 18\,500 \text{ km s}^{-1}$ ).

A wide-area observing campaign was therefore launched to assess the full extent of the Vela overdensity, using multiobject spectrographs on both the Southern African Large Telescope (SALT) and the 3.9-m Anglo–Australian Telescope (AAT). Section 2 will summarize the spectroscopic survey results, inclusive of the above-mentioned unpublished Optopus and 6dF redshifts, and Section 3 describes the resulting velocity distribution. Section 4 presents a discussion in support of a supercluster morphology and an assessment of the amplitude of the overdensity. The conclusions of this Letter and planned follow-up are discussed in Section 5.

## 2 SPECTROSCOPIC DATA

The spectroscopic survey focused on the ZOA region around Vela:  $245^\circ \lesssim l \lesssim 295^\circ$  and  $|b| \lesssim 10^\circ$ . The galaxy target sample was based on a combination of galaxies from deep optical galaxy surveys (Kraan-Korteweg & Lahav 2000; Kraan-Korteweg 2000), which had the specific aim of reducing the unmapped area in the southern Galactic ZOA, plus galaxies identified in the 2MASS Extended Source Catalogue (2MASX) (Jarrett 2000). The optical catalogue is a diameter-selected sample ( $D \geq 12 \text{ arcsec}$ ) which was found by Kraan-Korteweg (2000) to be complete to a magnitude limit  $B_J \leq 18.5\text{--}19.0 \text{ mag}$  up to extinction levels  $A_B = 3.0 \text{ mag}$ . The 2MASX finds galaxies closer to the Galactic plane (up to about  $A_B \lesssim 10 \text{ mag}$ ) but becomes less effective where the star density is high, and has a completeness limit of  $K_s \lesssim 13.5 \text{ mag}$ . The catalogues are complementary in the types of galaxies they uncover at low Galactic latitudes (Kraan-Korteweg & Jarrett 2005).



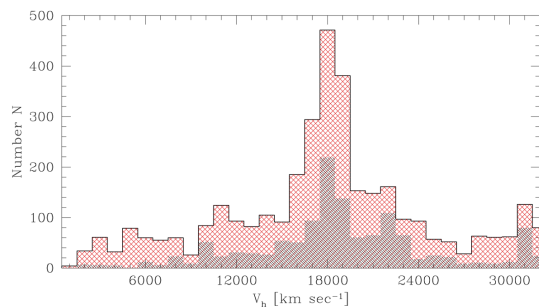
**Figure 1.** Redshift distribution in Galactic coordinates. Large dots mark new redshifts ( $N = 4432$ ), and crosses are previous redshifts. Colours indicate redshift (cyan:  $<8000$ ; green:  $8000\text{--}16\,000$ ; red:  $16\,000\text{--}20\,000$ ; magenta:  $20\,000\text{--}24\,000$ ; blue:  $24\,000\text{--}32\,000 \text{ km s}^{-1}$ ). Cyan circles mark the two 6dF fields (large) and the  $25\,2^\circ$  AAOmega fields (small). Contours indicate extinctions of  $A_K = 0.3 \text{ mag}$  (solid), above which the galaxy sample is incomplete, and  $0.18 \text{ mag}$  (dotted).

From 2012 to 2014, when the SALT multiobject spectrograph (MOS) became operational, time was allocated to observe Vela fields of particularly high galaxy concentrations (undersampled by 6dF observations). With its numerous slitlets over the relatively small field of view ( $8 \times 8 \text{ arcmin}^2$ ), the SALT MOS is ideal for observing cluster cores in this distance range. With typical exposure times of 12–20 min, reliable redshifts ( $\sigma_z \sim 150 \text{ km s}^{-1}$ ) were obtained for about 80 per cent of the potential cluster galaxies. The SALT observations identified 10 clusters (out of 13 targeted fields). This includes the two above-mentioned CIZA X-ray clusters, which had only one or two prior redshifts. Interestingly, 7 of the 10 clusters were found to reside in the redshift range of the Vela overdensity.

Major strides were achieved in 2014 using the 2dF+AAOmega spectrograph on the 3.9-m AAT, which has a  $2^\circ$  diameter field, in which up to 392 fibres can be positioned. We observed 25 fields with AAOmega over six nights. We employed a sparse-sampling strategy to cover as wide an area as possible in the allocated observing period; the locations of the fields are shown in Fig. 1. Most of the pointings are towards intermediate latitudes ( $|b| \sim 4^\circ\text{--}10^\circ$ ), where extinction is not excessive. Of the 4747 target galaxies in these 25 fields, 92 per cent could be fibred up. Exposure times were 60–90 min depending on foreground extinction and weather conditions. The success rate for science quality redshifts ( $\sigma_z \sim 100 \text{ km s}^{-1}$ ) of these partially obscured galaxies was 95 per cent. These AAOmega observations yielded over 4100 unique new redshifts in the ZOA.

Combining the AAOmega and SALT redshift data with our earlier unpublished data (Optopus, 6dF, and some 1.9-m spectroscopy from SAAO), a total of 4432 new redshifts were obtained (although 126 lie at  $|b| > 10^\circ$ ). These galaxies and their redshifts are plotted in Galactic coordinates in Fig. 1, including additional galaxy redshifts available from the literature. Note the scarcity of redshifts at  $|b| \leq 10^\circ$ , except at higher longitudes ( $l \gtrsim 280^\circ$ ) which our team surveyed in earlier work (Kraan-Korteweg et al. 1994). The cluster cores targeted with SALT (too small and dense to be outlined separately in the figure) mostly lie within the AAOmega fields, apart from one field at an extremely low latitude,  $(l, b) \sim (251.5, 0.9)$ , and two CIZA X-ray clusters (at  $b \sim -11.5$  and  $-12.5$ ).

Despite the sparse sampling within the latitude bands of intermediate extinction ( $|b| \gtrsim 4^\circ$ ), the predominance of galaxies in the redshift range of the Vela overdensity (red and magenta;  $16\text{--}24\,000 \text{ km s}^{-1}$ ) is prominent in 20 of the 25 widely spread AAOmega fields on both sides of the Galactic plane. The



**Figure 2.** Redshift distribution in the Vela region ( $255^\circ \leq l \leq 285^\circ$ ,  $|b| \leq 10^\circ$ ). Red shows all galaxies; grey shows galaxies below the Galactic plane.

overdensity appears to extend over most of our ZOA survey region, from about  $25^\circ$  above to  $20^\circ$  below the Galactic plane.

### 3 DISTRIBUTION IN REDSHIFT SPACE

To assess the significance of the overdensity, Fig. 2 shows the redshift distribution of all galaxies in the region  $255^\circ \leq l \leq 285^\circ$  and  $|b| < 10^\circ$ . There is a highly significant peak at  $17\,000$ – $19\,000$   $\text{km s}^{-1}$ , with broad shoulders extending over  $15\,000$ – $23\,000$   $\text{km s}^{-1}$ . Such an extent is typical of superclusters (Einasto et al. 2014), and, indeed, the velocity distribution is remarkably similar to that of the Shapley supercluster (SSC) obtained by Proust et al. (2006). We note that the Vela structure emerges at a distance  $cz \sim 16\,000$   $\text{km s}^{-1}$ , where existing ‘whole-sky’ redshift surveys lose sensitivity (Jones et al. 2009; Huchra et al. 2012; Tully, Courtois & Sorce 2016).

The grey-shaded area marks the redshift distribution of galaxies below the Galactic plane to emphasize that the overdensity is equally pronounced on either side of the Galactic plane, with the galaxies below the plane showing a secondary peak at  $cz \sim 22\,000$   $\text{km s}^{-1}$ . The similarity of the histogram suggests that the overdensity is contiguous across the plane, implying an overall extent of at least  $25 \times 20$   $\text{deg}^2$ , corresponding to  $\sim 115 \times 90$   $h_{70}$  Mpc.

Wedge diagrams are presented in Fig. 3 for all the galaxies delimited by the same survey area as Fig. 2; the right-hand panels show structures above (top) and below (bottom) the plane. The detail and depth of the large-scale structures unveiled in the ZOA are striking. The overdensity appears to consist of a main broad wall (W1) at

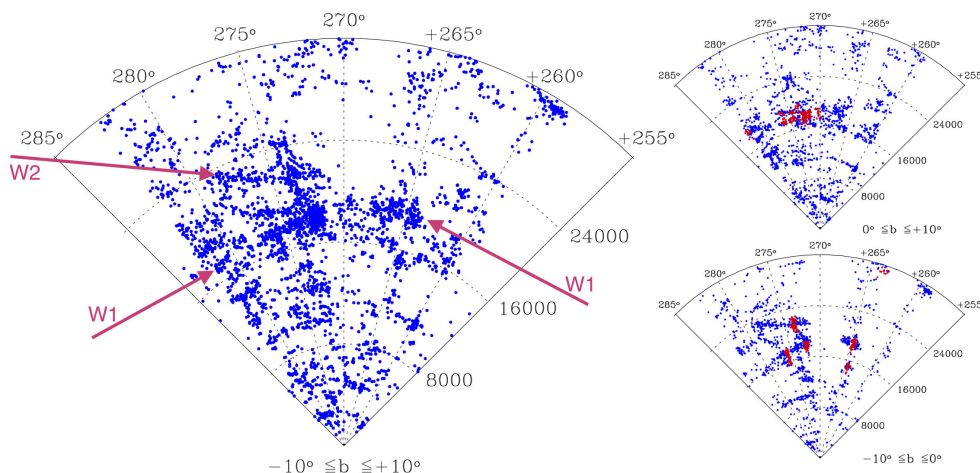
$cz \sim 18\,000$   $\text{km s}^{-1}$  and a smaller one (W2) at a slightly higher redshift ( $cz \sim 22\,000$   $\text{km s}^{-1}$ ; see Fig. 2). The main wall appears both above and below the plane, while the second wall is present only below the plane. Most of the gaps in these fairly smooth walls are the result of the sparse sampling to date (see Fig. 1). The main wall extends over at least  $120$   $h_{70}$  Mpc.

Both walls are quite broad, encompassing numerous clusters and groups. The SALT observations confirm seven clusters (five new) within these walls, with indications of further clusters in many of the AAOmega fields, despite undersampling due to fibre positioning constraints. At present the survey does not reach much below the characteristic magnitude at the mean Vela supercluster (VSCL) distance (see next section). A preliminary clustering analysis based on these redshift data, including an assessment of the velocity dispersion, cluster-centric distribution of redshifts, the spatial concentration, and the steepness of the extinction-corrected  $K_s^0$ -band luminosity function (LF), has led to the identification of 20 potential clusters with velocity dispersions  $\sigma_v > 400$   $\text{km s}^{-1}$  that could be part of the Vela overdensity (galaxies within an Abell radius are marked as red dots in Fig. 3).

### 4 VELA AS SUPERCLUSTER

*Supercluster morphology.* The spectroscopic data are consistent with the Vela overdensity being a supercluster: an elongated, prolate, wall-like structure (Einasto et al. 2011) with embedded clusters. Examinations of the walls from various perspectives, including photometric redshift data (2MPZ, Bilicki et al. 2014), suggest that the two walls below the plane merge in the obscured region, then emanate above the plane as one wall. This again is a typical morphology for a large supercluster, which can show multibranching of filaments and walls (Einasto et al. 2011). Both the observed cosmic web and simulations of structure formation (Springel et al. 2005) predict that massive clusters reside at the confluence of such walls. It is an unfortunate coincidence that the current optical data preclude charting the merger of these walls, which occurs in the most obscured part of the ZOA – we are most likely missing the core of the VSCL.

*Clusters.* An indicator of the mass of a supercluster is the number of its massive clusters. A total of 20 potential new clusters were identified in our data set (see Fig. 3), in addition to the two CIZA X-ray clusters just outside of our survey region. Surprisingly,



**Figure 3.** Left-hand panel: redshift cone out to  $cz \leq 32\,000$   $\text{km s}^{-1}$  for the full survey area ( $|b| < 10^\circ$ ). Right-hand panel: structures above ( $b > 10^\circ$ ; top panel) and below ( $b < 10^\circ$ ; bottom panel) the Galactic plane; red dots mark galaxies in potential clusters. The primary wall at  $18\,000$   $\text{km s}^{-1}$  is prominent on both sides of the plane, whereas the higher velocity wall evident below the plane seems to intersect the primary wall in the plane at  $l \leq 265^\circ$ .

Kocevski & Ebeling (2006) list no further CIZA clusters within the VSCL walls. Their catalogue is claimed to be fairly complete to Galactic gas column densities of  $N_{\text{HI}} \lesssim (3\text{--}5) \times 10^{21} \text{ cm}^{-2}$  ( $A_K \sim 0.18\text{--}0.3$  mag; i.e. outside the contours in Fig. 1). Above this, the low-energy X-ray bands of *ROSAT* (0.1–2.4 keV) are particularly susceptible to absorption by foreground gas – and the Galactic gas flares to higher latitudes ( $8^\circ\text{--}10^\circ$ ) over a large part of the Vela survey area (Fig. 1). To one side of the survey, in a region of diameter  $\sim 8.3$  centred on  $(l, b) = (263^\circ.9, -3^\circ.3)$ , the X-ray detection of clusters is extremely difficult due to the bright foreground emission from the Vela supernova remnant.

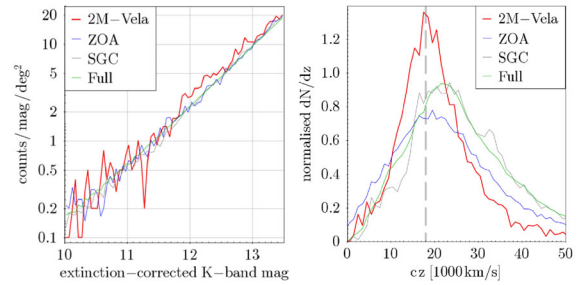
We explored the *ROSAT* All-Sky Survey (RASS) data base for X-ray emission around the position of the clusters identified in our redshift survey and detected X-ray emission for a few of them, even though the exposure of RASS is very shallow (200–700 s). The most prominent detection is a dynamically young cluster at  $cz = 18167 \text{ km s}^{-1}$  with a flux of  $F_X = 5.3 (\pm 0.7) \times 10^{-12} \text{ erg s}^{-1} \text{ cm}^{-2}$  and an X-ray luminosity of  $L_X = 4.6 \times 10^{43} \text{ erg s}^{-1}$  in the 0.1–2.4 keV band. With this luminosity, and with a complex distribution of emission, it resembles the Virgo cluster of galaxies.

The second brightest X-ray cluster has an X-ray luminosity  $L_X = 1.5 (\pm 0.4) \times 10^{43} \text{ erg s}^{-1}$ . Based on the precise X-ray LF determined in the REFLEX II and NORAS II surveys (Böhringer, Chon & Collins 2014), we can calculate the probability of finding clusters above this luminosity in our survey volume, taken generously as the VSCL region with  $N_{\text{HI}} \leq 3 \times 10^{21} \text{ cm}^{-2}$ . We find that about 0.8 detections should be expected; thus, the two detections we have are indicative of an overdensity.

**Overdensity.** One method to arrive at an estimate of the Vela mass overdensity, and its potential effect on the local velocity field, is through galaxy counts. To pursue this, we restricted the Vela survey sample to 2MASX galaxies and the survey region to latitude strips of  $6^\circ < |b| < 10^\circ$  (to minimize incompleteness biases) over the longitude range  $260^\circ < l < 285^\circ$  (within these strips, the Galactic extinction does not exceed  $A_K = 0.3$  mag for >99 per cent of the sources). We refer to this restricted sample henceforth as 2M-Vela. We inspected the apparent  $K_s$ -band counts and found them to be fairly complete to  $K_s = 13.8$  mag, though with a slight decline in counts fainter than 13.4 mag. We thus limit the Vela sample to  $K_s < 13.8$  mag, followed by a further cut to account for a maximum extinction of  $A_K = 0.3$  mag (assuming Schlafly & Finkbeiner 2011) – i.e.  $K_s^\circ < 13.5$  mag. This is close to the nominal completeness limit of 2MASX of  $K_s < 13.5$  mag away from the plane. The 2M-Vela sample allows us to estimate a lower limit to the galaxy overdensity.

For comparison, we defined three 2MASX subsamples with the same magnitude cut-offs. These are the full 2MASX galaxy sample (Full), limited to  $|b| > 15^\circ$  to avoid low-latitude incompleteness, the South Galactic Cap (SGC;  $b < -60^\circ$ ) and another ZOA sample (ZOA;  $110^\circ < l < 260^\circ$ ). The latter is restricted to the same latitude strips as 2M-Vela, and excludes the Galactic bulge; this six times larger ZOA area should be subject to similar incompleteness biases to 2M-Vela.

We first analysed the number counts per square degree (Fig. 4, left-hand panel). The  $\log N$  versus  $K^\circ$  relations for all these samples display the expected smooth linear increase up to the completeness limit (even the ZOA sample). However, 2M-Vela reveals, in addition, a clear elevation of counts for  $11.8 < K_s^\circ < 13.0$  mag (Fig. 4), exactly where an overdensity at  $18\,000 \text{ km s}^{-1}$  would reveal its signature. The characteristic magnitude derived by Kochanek et al. (2001) from the bright 2MASX catalogue,  $M_K^* = -24.2$  mag (adjusted to  $h_{70}$ ), corresponds to  $K_s^\circ < 12.8$  mag for the main VSCL wall (W1). Given the steepness of the LF at the bright end, the



**Figure 4.** The overdensity of 2M-Vela versus other 2MASX samples. Left-hand panel: direct number counts normalized to the survey areas. The VSCL is notable over the predicted 11.8–13.0 mag interval. Right-hand panel: photometric redshift distributions for the same samples, again normalized by area. The central velocity of the VSCL is indicated by the grey dashed line.

increase caused by the supercluster will drop to normal count levels at most one magnitude below  $M_K^*$  – as observed in Fig. 4. At the faint end, the increment will drop more rapidly, exacerbated by increasing incompleteness bias. There is no difference between the enhancement above and below the Galactic plane, except for the elevated counts remaining higher at the bright end (for about  $\Delta m \lesssim 0.2$  mag) below the plane, the second wall (W2).

We quantify the enhancement by dividing the 2M-Vela counts per  $\text{deg}^2$  by the respective numbers for the Two Micron All-Sky Survey (2MASS) comparison samples in the magnitude range in which the VSCL is detectable ( $11.8 < K_s^\circ < 13.0$  mag). This yields an enhancement in counts by a factor 1.17, 1.21, and 1.25 versus the Full, SGC, and ZOA samples, respectively.

In the second step, we estimate the Vela overdensity  $\delta$  ( $\delta \equiv \rho/\bar{\rho} - 1$ ) within a volume shell. We used the 2MPZ catalogue (Bilicki et al. 2014), which is a combination of spectroscopic redshifts (where available) supplemented by photometric ones. Caution in interpreting the data is warranted; large errors in photo- $z$  ( $\sigma_z \sim 0.015$ ) will smear out structures, making features less prominent and leading to lower overdensity estimates. In addition, the 2M-Vela and ZOA samples will be less complete because of their  $K_s^\circ = 13.5$  mag limit.

The right-hand panel of Fig. 4 shows redshift distributions, normalized by survey area, for the discussed samples. The peak at  $18\,000 \text{ km s}^{-1}$  in the Vela region is clearly visible. From these data, we determined the density of 2M-Vela and the three 2MASX comparison samples by subdividing them into redshift shells. For a shell of  $0.055 < z < 0.065$ , the resulting overdensity is  $\delta = 0.77$  when compared with the ZOA. As expected the values are slightly lower,  $\delta = 0.60$  and  $0.51$ , compared to 2MASS ( $|b| > 15^\circ$ ) or SGC, confirming that the ZOA samples are not as complete as the 2MASS and SGC samples. The values for wider shells are marginally lower (i.e. more diluted).

In summary, the VSCL is significantly overdense in 2MASS galaxy counts, and in a redshift shell centred at its average distance. The overdensity does not vary much on either side of the plane, giving further substance to the indications that the VSCL extends over at least  $25 \times 20 \text{ deg}^2$  on the sky.

Despite the limited sampling, we attempt an assessment of how VSCL compares to SSC. Extracting a sample in an area  $\sim 250 \text{ deg}^2$  around the SSC core in a shell of  $0.042 < z < 0.057$  from the 2MPZ leads to an overdensity of  $\delta_{\text{SSC}} \sim 1.4$ . This would, however, decrease 2.3-fold if the SSC were subjected to equivalent selection criteria to Vela (lower 2MASS completeness due to extinction and distance; dilution due to a high fraction of photometric redshifts), i.e.  $\delta_{\text{SSC}} \sim 0.6$ . We thus conclude that a similarly comprehensive

sampling of the VSCL would result in a similar overdensity to SSC, suggesting that the two structures may be comparable, particularly if the central core of VSCL is hidden behind the thickest dust layers.

*Implications for bulk flows.* The Vela overdensity estimate can be used for a crude assessment of its contribution to the LG velocity. A simple calculation based on linear perturbation theory gives  $v_{LG} \simeq 50 \text{ km s}^{-1}$ . This will likely be a lower limit due to the incomplete sampling. By comparison, the contribution from SSC is estimated – depending on the method – to be  $v_{LG} \simeq 55 \text{ km s}^{-1}$  (Loeb & Narayan 2008) or  $v_{LG} \simeq 90 \text{ km s}^{-1}$  (Lavaux & Hudson 2011). The influence of Vela on local bulk flows is thus probably comparable to SSC. A quantitative assessment will, however, require additional observations.

## 5 CONCLUSIONS

With close to 4500 new redshifts determined in the ZOA, in Vela, this preliminary data set provides clear evidence for the existence of a very extended ( $\sim 115 \times 90 h_{70}$  Mpc) overdensity at a redshift of  $\sim 18 \text{ 000 km s}^{-1}$ . Its morphology, with a broad main wall and a secondary merging wall, is characteristic of a supercluster. Twenty potential clusters were identified in the walls from our optical spectroscopy. This number constitutes a lower limit, as the surveyed fields cover only  $\sim 20$  per cent of the region, while the Galactic foreground gas limits the number of *ROSAT* X-ray cluster detections.

The Vela survey galaxy counts reveal an enhancement in counts of  $\sim 1.2$  compared to three 2MASX comparison samples over the magnitude range of the VSCL, while a redshift shell (based on spectroscopic and photometric redshifts from 2MPZ) shows excess densities of  $\delta = 0.50\text{--}0.77$  (depending on the comparison sample and the width of the shell). The current data set is as yet too sparsely sampled to reliably quantify the cosmological implications of the VSCL. However, preliminary calculations show that an overdensity of this magnitude would reduce the current misalignment of derived clustering dipoles and may play a role in accounting for the observed residual bulk flows (e.g. Scrimgeour et al. 2016).

A more comprehensive view of the VSCL requires systematic spectroscopic surveys to bridge the gaps between the current survey fields. The Taipan instrument (Kuehn et al. 2014) is ideal for wide-field follow-up spectroscopy, as it can target  $\sim 150$  objects per  $6^\circ$  field and will get reasonably complete redshifts, as faint as  $r = 17.5$  mag with 15-min integrations. Although the ZOA ( $|b| < 10^\circ$ ) does not form part of the Taipan survey, the VSCL, because of its potential science impact, has been selected for observation in the Taipan Science Verification phase in early 2017.

At higher extinctions, only systematic surveys in the 21-cm neutral hydrogen line will prevail (Staveley-Smith et al. 2016). We have proposed using the South African SKA Pathfinder MeerKAT in early science mode to survey the most obscured part of the VSCL. Our simulations have shown that this can be realized within reasonable time-scales with 32 dishes and does not need to await the full 64-dish array.

The VSCL is Terra Incognita, an unknown great continent in the nearby Universe whose outline we are only beginning to discern. We are pursuing an ambitious multiwavelength programme to cover the full width of the ZOA, including the opaque part, in order to uncover the full extent of the VSCL and determine its implications for cosmology.

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